

A search for OB stars in the field of the galactic OB association Bochum 7[★]

I. *UBV*-IR photometry and optical spectroscopy

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ABSTRACT

Context. The presence of a Wolf-Rayet star, WR12, at $\alpha = 8^{\text{h}}44^{\text{m}}47.2^{\text{s}}$, $\delta = -45^{\circ}58'55.5''$ (J2000.0), triggered the search for a surrounding OB association now known as Bochum 7. Several papers have dealt with the existence of Bo 7, but none actually included spectral types of the stars. To assure the spectral types of members of the association, it is essential to have spectral images of them.

Aims. In this paper we present previously unpublished spectra of OB stars over a 30 arcmin field covering the estimated field of the galactic OB association.

Methods. From *UBV* aperture photometry, OB star candidates were selected, and optical spectroscopic observations were obtained for these in CASLEO, Argentina.

Results. 63 stars were classified as OB type, including 8 stars already present in the LSS catalogue, and 37 of these 63 stars did not have a well-known spectral type. The OB stars were used to test the properties of the interstellar dust in the line of sight and the relation between the diffuse interstellar band at 4428 Å, which grows linearly with $E_{(B-V)}$ until 1 mag. Combining our data with 2MASS IR photometry we tested and confirmed that derivation of stellar parameters from our photometry, using the CHORIZOS code, is highly reliable.

Key words. Galaxy: open clusters and associations: individual: Bochum 7 – stars: early-type – techniques: spectroscopic – techniques: photometric

1. Introduction

Relatively low extinction at the region in the third quadrant in the Galactic plane allows deep analysis of the stellar population and galactic structure towards the outskirts of our Galaxy. The lack of tightly bound and easily identifiable young open clusters leads to the quest for OB associations as tracers of spiral arms. Massive stars are usually found in groups, and the presence of a Wolf-Rayet star, WR12 in the catalogue of galactic WR stars (van der Hucht 2001), at $\alpha = 8^{\text{h}}44^{\text{m}}47.2^{\text{s}}$, $\delta = -45^{\circ}58'55.5''$ (J2000.0), triggered the search for a surrounding OB association. Moffat & Vogt (1975) suggested that several OB stars listed in the LSS catalogue (Stephenson & Sanduleak 1971) around WR12 may be an OB association, namely Bochum 7, at 5.8 kpc, although the membership of WR12 itself was in doubt. Lundstrom & Stenholm (1984) proposed that Bo 7 is an OB association of ~ 30 pc of size. On the other hand, Sung et al. (1999) suggested that the stars associated with Bo 7 are only a local concentration in the density of young stars belonging to the Vel OB3 association. From data of the 7 stars listed in the LSS catalogue, Corti et al. (2003) derive a spectroscopic and kinematical distance of about 5.0 kpc for Bo 7, supporting the idea

that the Wolf-Rayet star LS 1145 = WR 12 was also member of this association.

In this work, we present the analysis made with photometric and spectroscopic data of the stars in a 30 arcmin field surrounding WR 12. With these we obtained spectral types, colour excesses, equivalent widths of the DIB 4428 feature, and distances for the OB stars found in the observed field. By comparing stellar temperatures and reddening derived from photometric and spectroscopic data we analyse the reliability of the (more abundant) photometric ones.

The paper is organized as follows: in Sect. 2 we describe the observations. In Sect. 3 we present the results and their discussion, and in Sect. 4 we summarize the main issues. In a later paper (Paper II) we will increase the number of stars that will be analysed, including OB stars without spectral data. We will estimate the distances and colour excesses of other possible members. Combining this information with proper motions we will be able to identify the stellar groups and to infer their evolutionary stages.

2. Observations

2.1. Photometric data

We obtained *UBV* photometric images with the 1-m Curtis-Schmidt telescope and a Thompson 1024 × 1024 CCD

[★] Figures 6–15 are only available in electronic form at <http://www.aanda.org>

with 19 μm pixels at the Cerro Tololo Inter-American Observatory (CTIO), Chile, in March 5, 1994. This instrumental configuration yields a 30×30 arcmin field of view, centred at $\alpha = 8^{\text{h}}44^{\text{m}}23^{\text{s}}$, $\delta = -45^{\circ}56'22''$ (J2000.0) ($l = 265^{\circ}.12$, $b = -2^{\circ}$). The scale on the detector is 1.834 arcsec/pixel, which subsamples a stellar image under average seeing conditions at CTIO. CCD acquisition was set to obtain images with $\text{GAIN} = 3.3 \text{ e}^-/\text{ADU}$ and $\text{RDNOISE} = 7\text{e}^-$.

Different exposure times were used for each filter: 10 s for the V and B filters and 20 s for the U filter. The same number of images were taken for the Landolt photometric standard region sa104 (Landolt 1992). For these images the exposure times were 10, 20, and 60 s for the V , B , and U filters respectively. The transformation relations used here were:

$$\text{fit } u1 = 0.0, u3 = 0.000$$

$$\text{const } u4 = -0.022, u2 = 0.591, u5 = -0.023$$

$$mU = (UB + BV + V) + u1 + u2 \cdot XU + u3 \cdot UB + u4 \cdot UB \cdot XU + u5 \cdot BV \cdot XU \quad (1)$$

$$\text{fit } b1 = 0.0, b3 = 0.000$$

$$\text{const } b4 = -0.023, b2 = 0.276$$

$$mB = (BV + V) + b1 + b2 \cdot XB + b3 \cdot BV + b4 \cdot BV \cdot XB \quad (2)$$

$$\text{fit } v1 = 0.0, v3 = 0.000$$

$$\text{const } v4 = 0.007, v2 = 0.152$$

$$mV = V + v1 + v2 \cdot XV + v3 \cdot BV + v4 \cdot BV \cdot XV, \quad (3)$$

where mU , mB , and mV are the instrumental magnitudes; UB and BV are the colour indices; V is the visual magnitude; and XU , XB , and XV are the air masses of each filter. We also downloaded near-infrared JHK magnitudes for stars within our observed area from the Two Micron All Sky Survey (2MASS) Point Source Catalog (Cutri et al. 2003).

2.2. Digital spectra

Spectral images were obtained at the Complejo Astronómico El Leoncito (CASLEO)¹, San Juan, Argentina, during several observing runs between 2000 and 2002. The instrumental configurations used were: the Boller and Chivens (B&C) and the REOSC Cassegrain spectrographs attached to the 2.15-m telescope at CASLEO. In both spectrographs we used a 600 l mm^{-1} grating, and the slit width was set to match an aperture of 2 to 3 arcsec according to the average seeing at CASLEO. Fifteen spectra were secured with the B&C spectrograph during January 2001. This spectrograph holds a Photometrics CCD of 512×512 pixels, at a $20 \mu\text{m}$ pixel size. The reciprocal dispersion of these spectra is $\approx 2.3 \text{ \AA px}^{-1}$ and the spectral range 3800–5000 \AA . The signal-to-noise ratio was always higher than 120. One hundred and eighty spectra were obtained with the REOSC Cassegrain spectrograph during January 2000, February 2001, and January and February 2002,

using a TEK 1024×1024 pixel CCD detector, with a pixel size of $24 \mu\text{m}$. The reciprocal dispersion was $\approx 1.65 \text{ \AA px}^{-1}$. The spectral range of these spectra is 3700–5200 \AA and the signal-to-noise ratio ranges from 150 to 200.

He-Ar comparison lamp images were observed at the same telescope position as the stellar images immediately after or before the stellar exposures. Bias frames were also obtained every night, as well as spectra of the planetary nebula IC418 and HR 1996 as radial velocity standards stars. All spectra were processed and analysed with IRAF² routines at La Plata Observatory.

3. Results and their discussion

3.1. Photometry

IRAF's standard image reduction procedures were followed and aperture photometry was calculated on science and standard images, as subsampling prevented us from performing point-spread function photometry. The aperture radius chosen was 1.5 pixels, and the sky annulus inner radius was 5 pixels from the star centre and 3 pixels wide. Approximately 2000 stars were identified in the field of view including stars as faint as $V \sim 18$.

We compared our photometric results with those of Sung et al. (1999), who had obtained $UBVI$ CCD photometry for the region centred at $\alpha = 8^{\text{h}}44^{\text{m}}47.2^{\text{s}}$, $\delta = -45^{\circ}58'9''$ (J2000.0) using the 1-m telescope (f/8) at Siding Spring Observatory. To do so, we first derived an astrometric calibration for our stars, using the ASTROM task within STARLINK software package. We then correlated our photometry table with that of Sung et al. (1999), matching stars by their celestial coordinates. In this way, we found ≈ 1100 stars in common, with more than 300 stars with magnitudes brighter than 15 in the V -band. The comparison between both photometries can be seen in Fig. 2, where it is possible to detect a minor systematic difference (0.04 mag) between our photometry and that of Sung et al. (1999). We have also used their photometry as reference for estimating an upper limit for the overall uncertainty in our derived magnitudes, obtaining the magnitude differences between our photometry and that of Sung et al. (1999), which was found to be about 0.08 mag in V , 0.10 mag in B , and 0.18 mag in U . This way, the uncertainty in our colours were 0.13 mag in $(B - V)$ and 0.21 mag in $(U - B)$. The colours $(B - V)$ and $(U - B)$ and the magnitudes V are listed in Table 1.

We analysed our UBV photometric data and selected more than 100 stars brighter than $V = 15$ with colours belonging to early-type stars. The selection was done using the reddening-free parameter Q (Schmidt-Kaler 1982). Although it would be safer to pick only stars with $Q < -0.5$ as early-type candidates, the number of stars with Q satisfying this and also brighter than 15 mag was too small. We therefore included stars with $-0.5 < Q < -0.3$, although the probability of detecting true B-type stars was unavoidably reduced. This brightness limit was included to build our target list for the spectroscopic runs described in Sect. 2.2 as the need of high signal-to-noise spectra imposes further restrictions on the target magnitudes for obtaining spectra at CASLEO.

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² IRAF is distributed by NOAO, operated by AURA, Inc. under an agreement with NSF.

Table 1. Photometric and spectroscopic information for individual stars analysed in this work.

ID	V	$(B - V)$	$(U - B)$	SpT	eqw4428	$V_o - M_o$	Cross ref.
CBN84332.2-455337	10.1	0.30	-0.41	B2IV	0.60	11.1	LS1132
CBN84433.4-454115	10.4	0.23	-0.47	B2III	0.85	12.5	2MASS 8443341-4541167
CBN84501.9-455034	10.4	0.18	-0.45	B2IV	0.73	11.9	CD-454491
CBN84446.2-460433	10.7	0.08	-0.49	B3V	0.60	11.2	CD-454481
CBN84324.0-460831	10.8	0.50	-0.50	O7.5V	1.90	12.8	LS1131 [▲]
CBN84447.2-455856	10.8	0.51	-0.36	WN8 [◊]	LS1145
CBN84506.0-455008	10.9	0.25	-0.35	B2V	0.60	11.3	CD-454493
CBN84349.6-460711	10.9	0.34	-0.61	O6.5V((f))	1.60	13.5	LS1135 [◊]
CBN84447.4-460339	11.1	0.09	-0.39	B3-5V	0.70	11.4	CD-454483
CBN84510.4-455854	11.2	0.55	-0.48	O9.5II	2.50	13.8	CPD-452977
CBN84504.8-460013	11.2	0.14	-0.44	B2V	0.70	12.2	CD-454492
CBN84406.6-455642	11.3	0.17	-0.32	B2V	0.80	12.2	CD-454466?
CBN84424.9-455334	11.3	0.61	-0.44	O7.5V	1.90	12.9	LS1144 [▲]
CBN84426.0-455749	11.4	0.25	-0.16	B8V	0.60	10.3	CD-454473
CBN84450.1-455648	11.4	0.16	-0.51	B2V	0.80	12.3	CD-454484
CBN84350.9-460348	11.4	0.45	-0.49	O9-9.5V	1.90	13.0	LS1137
CBN84506.3-460906	11.5	0.55	-0.47	B0IIIe	1.90	12.5:	LS1147 [▲]
CBN84543.1-460742	11.5	0.10	-0.42	B2-3V	0.40	12.4	CPD-453000
CBN84450.9-461012	11.6	0.47	-0.48	B1III	2.38	13.4	LS1146 [▲]
CBN84404.4-455316	11.7	0.68	-0.33	B0III	2.27	13.2	LS1140
CBN84339.5-455148	11.7	0.54	-0.14	B2III	0.74	12.9	CPD-452909
CBN84518.6-454543	11.8	0.27	-0.30	B2V	0.48	12.4	CPD-452984
CBN84325.3-460750	11.8	0.25	-0.05	B8-A0V	0.40	10.4	CPD-452901
CBN84552.3-460929	11.8	0.10	-0.41	B5V	0.44	12.1	
CBN84359.6-460237	12.0	0.20	-0.28	B5-8V	0.70	11.5	CPD-452926
CBN84338.5-460817	12.0	0.45	-0.52	O9.5V	1.90	13.5	CPD-452910
CBN84349.4-454204	12.0	0.29	-0.25	B3-5V	...	11.7	CPD-452917
CBN84443.0-455130	12.0	0.21	-0.28	B8V	0.50	11.0	CPD-452953
CBN84318.7-460742	12.2	0.24	-0.19	B8V	...	11.1	CPD-452895
CBN84438.9-460746	12.2	0.10	-0.36	B5V	0.50	12.4	2MASS 8443908-4607450
CBN84327.7-460929	12.3	0.63	-0.58	B0V-III	2.20	13.4	2MASS 8432796-4609263
CBN84348.6-460736	12.3	0.40	-0.47	B1-1.5V	1.70	12.9	2MASS 8434879-4607337
CBN84434.3-455214	12.5	0.23	-0.30	B5V	1.40	12.3	2MASS 8443427-4552136
CBN84505.0-460300	12.6	0.16	-0.37	B2.5-3V	0.40	12.9	2MASS 8450512-4602598
CBN84446.7-454632	12.6	0.46	-0.09	B3-5V	0.40	11.8	2MASS 8444674-4546333
CBN84327.4-460240	12.6	0.56	-0.44	B0V	1.60	13.4	2MASS 8432763-4602382
CBN84500.4-455834	12.6	0.24	-0.34	B5V	0.70	12.3	2MASS 8450044-4558343
CBN84505.7-454040	12.7	0.60	-0.31	B1III	2.60	14.0	2MASS 8450573-4540420
CBN84526.1-460527	12.9	0.66	-0.09	B2III	2.10	13.8	2MASS 8452620-4605258
CBN84350.7-455820	12.9	0.55	-0.34	B1.5V	2.60	12.5	2MASS 8435084-4558194
CBN84344.7-460656	12.9	0.43	-0.44	B1V	1.90	13.5	2MASS 8434493-4606537
CBN84522.4-455527	13.0	0.72	-0.24	B1III	2.60	13.9	2MASS 8452244-4555280
CBN84502.0-455257	13.2	0.23	-0.11	B8-A0V	0.70	11.8	2MASS 8450219-4552572
CBN84410.6-455452	13.3	0.37	-0.08	B5-8V	0.80	12.2	2MASS 8441074-4554522
CBN84419.9-455552	13.5	0.65	-0.08	B2III	2.90	14.3	2MASS 8441999-4555524
CBN84458.5-455242	13.5	0.24	-0.11	B8-9V	...	12.3	2MASS 8445857-4552434
CBN84431.5-455314	13.5	0.61	-0.34	B2III:	2.36:	14.3:	2MASS 8443163-4553148
CBN84518.2-455816	13.6	0.65	-0.32	B0.5V	3.00	14.1	2MASS 8451822-4558166
CBN84423.1-460821	13.7	0.65	-0.33	B1-2/V-III	1.80	14.0	
CBN84325.6-460820	13.7	0.62	-0.34	B2Ve	2.00	14.5:	2MASS 8432574-4608174
CBN84421.9-455136	13.7	0.66	-0.24	B1V	3.48	13.9	
CBN84437.8-455257	13.8	0.43	-0.32	B8-A0V	0.30	11.8	2MASS 8443788-4555288
CBN84510.9-455900	14.0	0.61	-0.31	B1V	2.32	14.2	
CBN84442.0-460802	14.1	0.51	-0.14	B2V	1.10	14.0	2MASS 8444210-4608004
CBN84346.7-460641	14.2	0.37	-0.43	B1-5V	2.20	14.2	2MASS 8434686-4606390
CBN84430.9-455442	14.2	0.53	-0.30	B2V	1.77	14.1	
CBN84523.5-455340	14.3	0.51	-0.06	B8V	1.80	12.4	2MASS 8452359-4553411
CBN84416.7-461121	14.3	0.53	-0.25	B1.5V	1.90	14.2	2MASS 8441683-4611189

3.2. Spectroscopy

3.2.1. Spectral types

The spectral classification of selected stars was done by comparing obtained spectra with the Digital Atlas of OB Stars

published by Walborn & Fitzpatrick(1990, WF90). As the spectral resolution of our data (1.65 \AA px^{-1}) was lower than the WF90 atlas (0.5 \AA px^{-1}), we have also obtained high signal-to-noise spectra ($S/N \geq 200$) of the atlas stars to detect important features for the classification of early-type

Table 1. continued.

ID	V	(B - V)	(U - B)	SpT	eqw4428	$V_o - M_o$	Cross ref.
CBN84423.2-455337	14.5	0.58	-0.24	B1.5-2V	1.15	14.1	2MASS 8442335-4553371
CBN84418.2-455307	14.6	0.56	-0.32	B1V	1.30	14.7	2MASS 8441838-4553078
CBN84414.0-455733	14.7	0.76	-0.16	B1.5-2V	1.90	13.9	2MASS 8441413-4557329
CBN84421.9-460125	14.8	0.49	-0.27	B1V	2.00	15.2	2MASS 8442208-4610224
CBN84518.8-454147	14.8	0.57	-0.13	B5V	0.87	13.3	2MASS 8451888-4541494
CBN84527.2-455360	14.9	0.49	+0.01	B5-8V	1.14	13.5	2MASS 8452725-4554006

* van der Hucht (2001).

Column 1 lists the stars' identification according to the guidelines of the IAU, i.e., using the (capitalized) first letters of the last names of the authors and the numbers, of the celestial coordinates in the J2000.0 frame. Columns 2 to 4 list the *UBV* magnitude and colours; the MK spectral classification is listed in Col. 5. The $\lambda 4428$ equivalent width is showed in \AA in Col. 6. Column 7 shows the spectrophotometric distance modulus determined adopting the absolute magnitudes corresponding to the spectral types, according to the calibration by Martins et al. (2005) for O spectral types and Schmidt-Kaler (1982) for B spectral types. Column 8 includes the cross reference for these stars, where LS indicates stars in the Luminous Star Catalogue (Stephenson & Sanduleak 1971), CD indicates stars in the Córdoba Durchmusterung catalogue, CPD indicates stars in the Cape Photographic Durchmusterung catalogue, and 2MASS indicates stars in the 2MASS catalogue (Cutri et al. 2003).

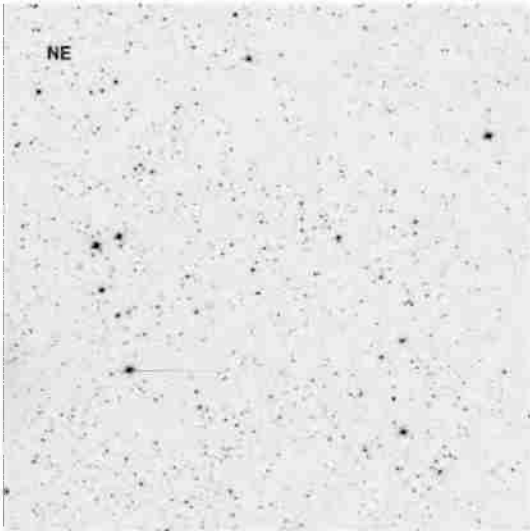


Fig. 1. V frame of Bochum 7 centred at $\alpha = 8^{\text{h}}44^{\text{m}}47.2^{\text{s}}$, $\delta = -45^{\circ}58'9$ (J2000.0). The scale on the detector is 1.834 arcsec/pixels and the field covered is 30 arcmin on each side.

stars. However, the signal-to-noise ratio decreases for fainter stars, leading to somewhat larger uncertainties in the spectral classification. With these data we are reproducing the WF90 atlas for lower dispersion spectrographs. A preliminary version of the atlas is already available on the internet at <http://fcaglp.unlp.edu.ar/~mariela/OBSpectralAtlas>.

From a final set of 102 stars, we classified 63 stars as early (OB) type, among which 50 are main sequence objects and 13 are subgiants, giants, or luminous giants. Among these 63 OB type stars, we found new spectral types for 37 stars. From Fig. 4 to Fig. 15 show the complete sample of classified spectra. The complete list can be found in Table 1. The availability of our Atlas of OB standards allowed us to reassess previous spectral classification (Corti et al. 2003). In this way, we classified LS 1132 as a B2IV and LS 1147 as a B0IIIe, previously classified as B3V and B1Ve, respectively. We searched the data about the spectral types of these 63 stars in the ‘‘Catalog of Galactic OB Stars’’ (Reed 2003), and we only found our own references (Corti et al. 2003) for the stars in the Luminous Star Catalogue (Stephenson & Sanduleak 1971), together with SpT = O6III (Vijapurkar & Drilling 1993) and O7Vn (Vijapurkar & Drilling 1993) for LS 1135 and LS 1131, respectively.

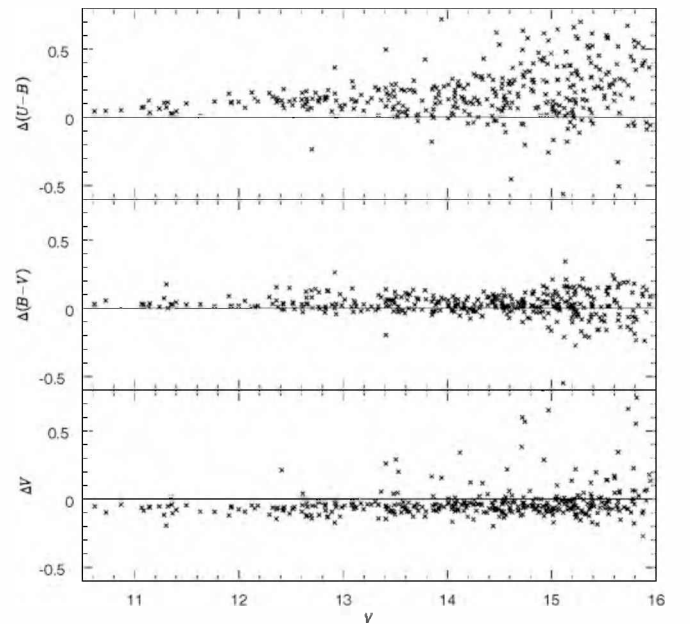


Fig. 2. Differences between our photometry and that of Sung et al. (1999). There is a minor difference between both photometries of about 0.04 mag in V. The lines for 0 mag and 0 colours show the offsets more clearly.

With our spectral classification we are able to test the derivation of basic stellar parameters, such as temperature and observed extinction, from photometric data. This will be useful when analysing the larger sample of stars for which we only have photometry available for the *UBVJHK_s* analysis and for which we employed the CHORIZOS code (Maíz-Apellániz 2004). For this we used the atmospheres of Kurucz (2004) with $Z = 0.0$ and main sequence gravity as model atmospheres. The Kurucz atmospheres were extinguished using $R_{5495} = 2.0-6.0 + \text{MC}$ laws and $E(4405-5495)$ ranging from -0.5 to 5.0 . The CHORIZOS code did the zero point (hereafter ZP) correction in all the synthetic magnitudes, employing the ZP correction of Maíz Apellániz (2006) for the Johnson system *UBV* and the ZP correction of Cohen et al. (2003) for the 2MASS *JHK_s* photometry. Figure 3 shows the comparison of the star's temperature obtained with the spectral classification and that obtained with the photometric data. An overall good agreement, within observational errors, can be readily see. Stars with notorious discrepancies are labelled in the plot and discussed below. Among the hotter stars,

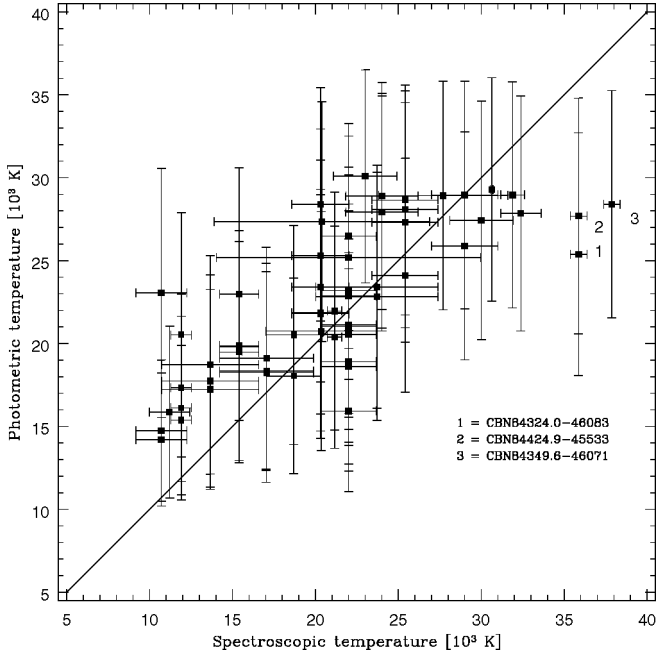


Fig. 3. Relation between the temperature values obtained from our photometric data using the CHORIZOS code (y -axis) and the data derived from known spectral types (x -axis). CBN84349.6-460711 is a binary system and CBN84324.0-460831 and CBN84424.9-455334 are two real candidates of binary systems. See discussion in the text.

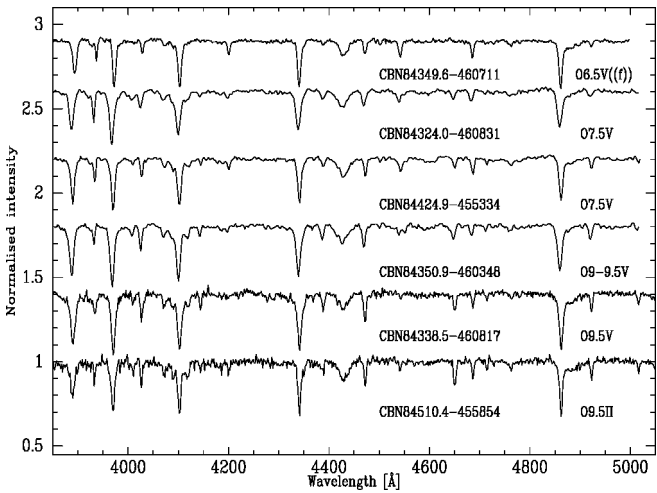


Fig. 4. Sample spectra of stars in the region of Bochum 7. Spectra are normalised. This figure includes a wide range of O spectral types. The complete list can be found in Table 1.

CBN84349.6-460711 is a binary system (Corti et al. 2003) and CBN84324.0-460831 and CBN84424.9-455334 are two strong candidates of binary system (Corti 2005). The last two stars show large variations among the radial velocities at the absorption lines in their respective spectra (Corti, in preparation). The reason for the discrepancy between the temperature obtained with the spectral classification and photometric data of these stars could be that the secondary component participates in their colours and for this reason the temperature obtained with the photometric data does not agree with the spectroscopic temperature.

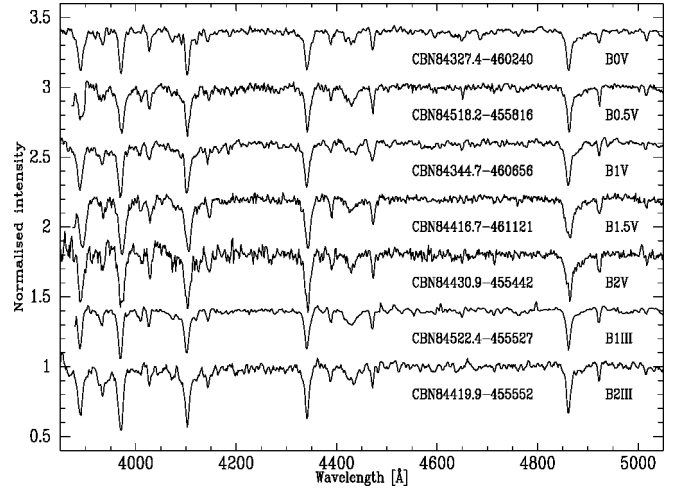


Fig. 5. Sample spectra of stars in the region of Bochum 7. Spectra are normalised. This figure includes a wide range of B spectral types. The complete list can be found in Table 1.

3.2.2. 4428 Å diffuse interstellar band (DIB)

We studied the 4428 Å DIB in our medium resolution spectra of OB stars. We fitted a Gaussian function to the profile of the DIB in 6 O-type and 54 B-type stars of our sample, and measured the equivalent width (EW) of the DIB in the spectrum of each star. EWs are listed in Table 1. This DIB is well evident in the blue optical spectra in Fig. 4. Stars of early spectral type, such as those included in this work, are especially useful for 4428 Å diffuse band studies.

Figure 16 shows the correlation of $\lambda 4428$ equivalent width with the stellar colour excess $E_{(B-V)}$, for the sample shown in Fig. 1. We have added data points from Herbig (1975) so as to include DIB measurements for stars of the same reddening as those in our sample. We have also added points from Snow et al. (2002) representing very heavily reddened lines of sight in Cygnus. The plot shows a positive correlation, as suggested by Herbig (1975), at least up to roughly $E_{(B-V)} \approx 1.0$ mag in our data. However, we find a steeper slope for the linear fit through our data than Herbig (1975), and it can also be seen that our data seem to complement those of Snow et al. (2002) better.

3.3. Interstellar extinction

The combination of photometric and spectroscopic information provides the only chance of determining individual colour excesses of the program stars, and hence of analysing the behaviour of the interstellar extinction in the line of sight. Figure 17 shows the photometric $E_{(B-V)}$ obtained using the CHORIZOS code with the atmosphere models mentioned previously vs. spectroscopic $E_{(B-V)}$, both with their error bars. It is possible to see a very good correspondence between both.

In the colour–colour diagram (Fig. 18), we have plotted the sixty-three OB stars and the other thirty-nine stars with spectral types between A0 and G3. It can be seen that the contamination of stars earlier than B8 type increases sharply as we move down to A0 type stars (Schmidt-Kaler 1982). There are two stars named CBN84431.8-455813 and CBN84446.0-45544 that do not fall where expected according to their spectral types. We reviewed their photometry and their spectral data and could not find any visible source for this odd behaviour, although they could be unresolved visual binaries.

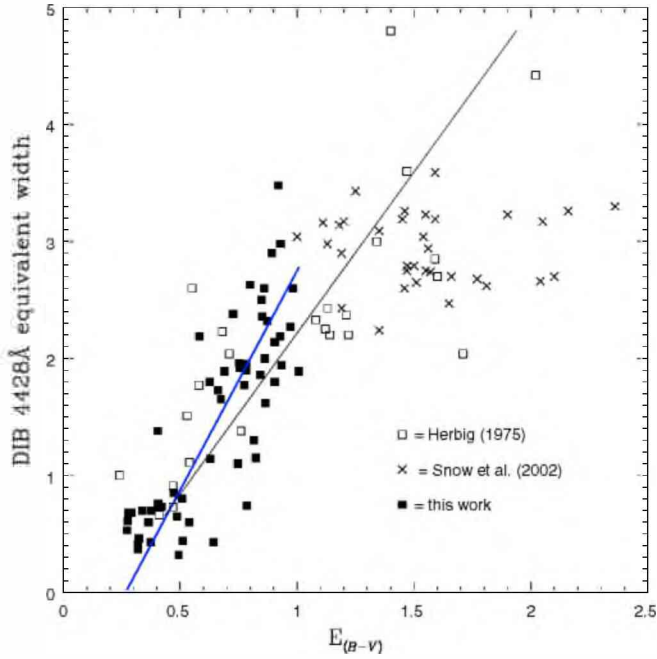


Fig. 16. Correlation of $\lambda 4428$ equivalent width with colour excess $E_{(B-V)}$. The open squares are data from Herbig (1975), the crosses are data from Snow et al. (2002), and the full squares are our own data. Within the broad scatter evident in this diagram, we see that the equivalent width of the 4428 Å DIB levels off with increasing extinction beyond $E_{(B-V)} \simeq 1.0$ mag.

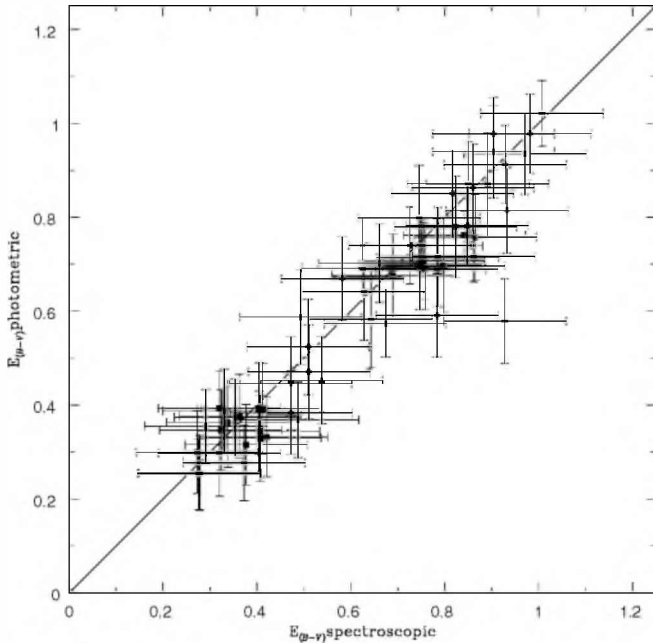


Fig. 17. Comparison of the colour excesses ($B - V$) obtained using the combination of photometric and spectroscopic information of our stars in the x -axis and using only photometric with CHORIZOS code (Maíz-Apellániz 2004) in the y -axis. Both data sets have their respective errors.

3.4. Distance

In a previous paper (Corti et al. 2003) a mean distance modulus for the LS stars in the Bo 7 area was obtained using Eq. (4), where M_V (Schmidt-Kaler 1982) was estimated from the stars' spectral type and A_V was directly calculated from their intrinsic

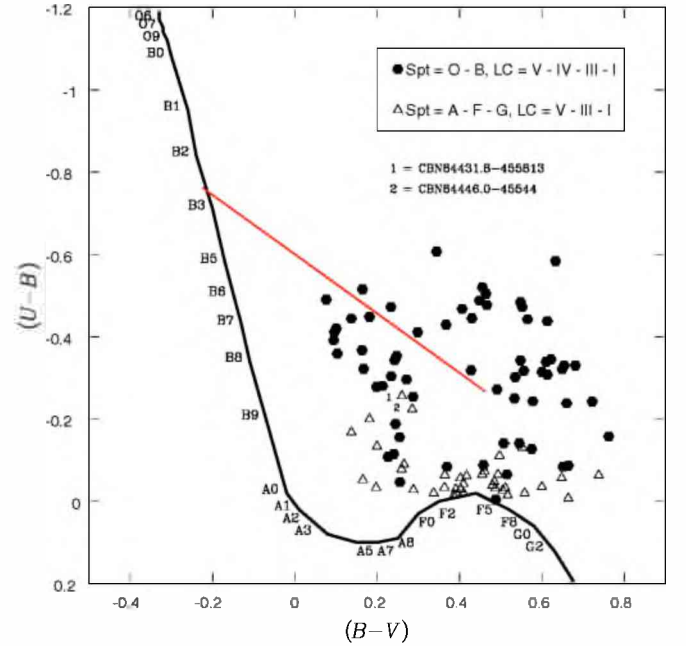


Fig. 18. Colour-colour diagram of the stars in the field of Bochum 7 with photometric and spectroscopic data. The sixty-three OB stars are plotted with filled hexagons and the other thirty-nine stars with spectral types between A0 and G3 are plotted with open triangles. The line represents the standard extinction curve for the B3 spectral types. The stars named CBN84431.8-455813 and CBN84446.0-45544 do not fall where expected according to their spectral types. They could be unresolved visual binaries.

sic colours (Schmidt-Kaler 1982), assuming a canonical relation between total and selective extinction. We can now extend the analysis to our larger sample and calculate the distance for the 63 OB spectral type stars. For these we employed the value of R obtained by using the CHORIZOS code, mentioned previously, for each star. For those stars with emission lines, where the R determination is unreliable, we replaced that value with $R = 3.7 \pm 0.3$, obtained averaging all values of R resultants for each star of the CHORIZOS code.

$$V - M_V = 5 \times \log_{10}(d) - 5 + A_V. \quad (4)$$

The value of mean distance obtained for the 63 stars is 4.2 ± 2.1 kpc. The derived distances for the 63 stars are listed in Table 1. The analysis of this information combined with the kinematic of the proper motion, will be discussed in Paper II, in which we determine the membership to stellar groups.

4. Summary

Doing *UBV* and *IR* photometry and optical spectroscopy of the 63 OB spectral type stars found towards Bochum 7, we have:

- Obtained 6 stars with spectral type O and 57 stars with spectral type B. Among these 63 stars, there are 11 giants and 52 dwarfs of main-sequence (MS) stars; 6 giants and 31 dwarfs of MS stars were discovered by us.
- Tested and confirmed that derivation of stellar parameters from our photometry (using CHORIZOS code) is highly reliable.
- Analysed the behaviour of interstellar extinction, finding that the spectral feature of the diffuse interstellar band 4428 Å grows linearly with $E_{(B-V)}$ until 1 mag. For the rest

of values of $E_{(B-V)}$, the feature of the DIB 4428 Å is approximately constant.

- Estimated spectrophotometric distances to 63 stars, lying between 1.2 and 10.8 kpc from the Sun.

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Online Material

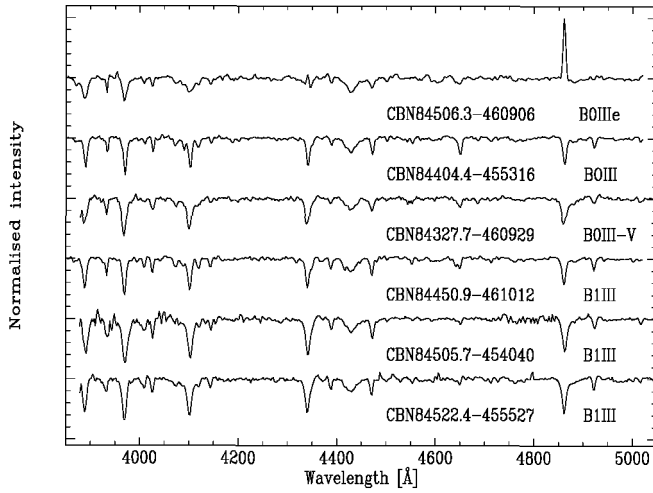


Fig. 6. This figure includes stars with B0-III and B0IIIe spectral types in the region of Bochum 7. 1.

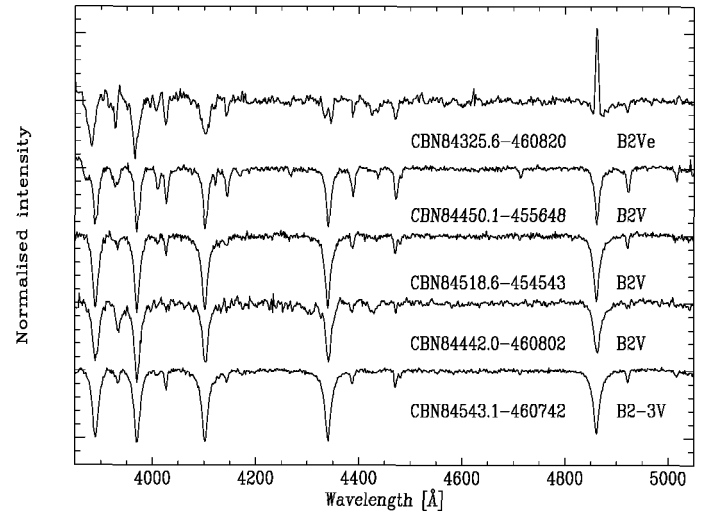


Fig. 9. This figure includes stars with B2-3V and B2Ve spectral types in the region of Bochum 7. 1.

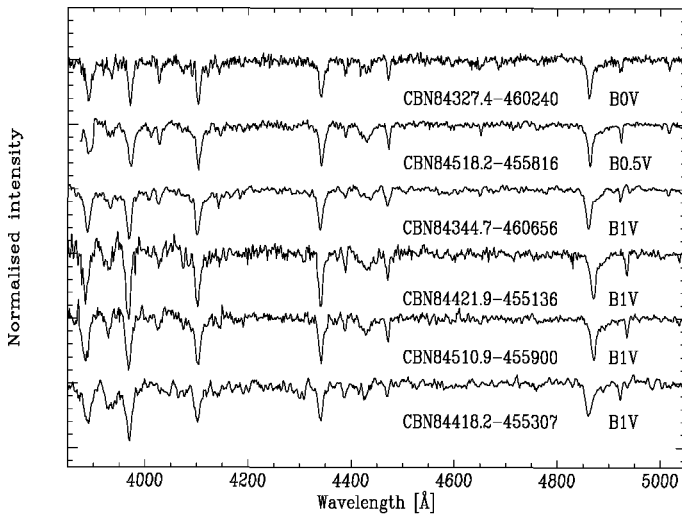


Fig. 7. This figure includes stars with B0-1V spectral types in the region of Bochum 7. 1.

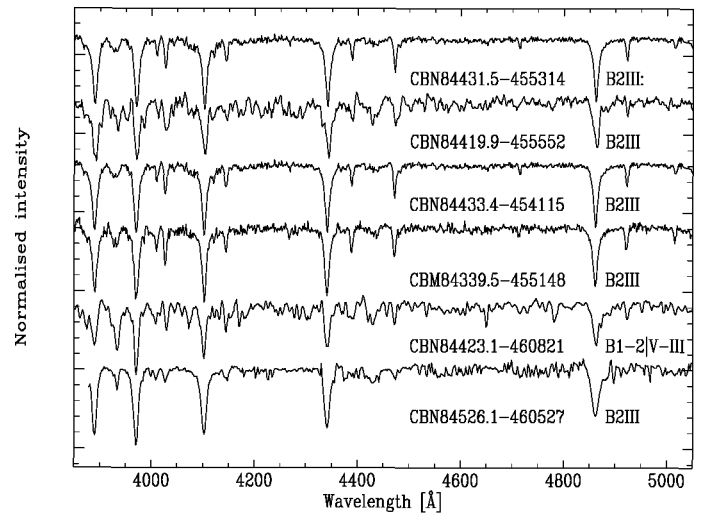


Fig. 10. This figure includes stars with B2III spectral types in the region of Bochum 7. 1.

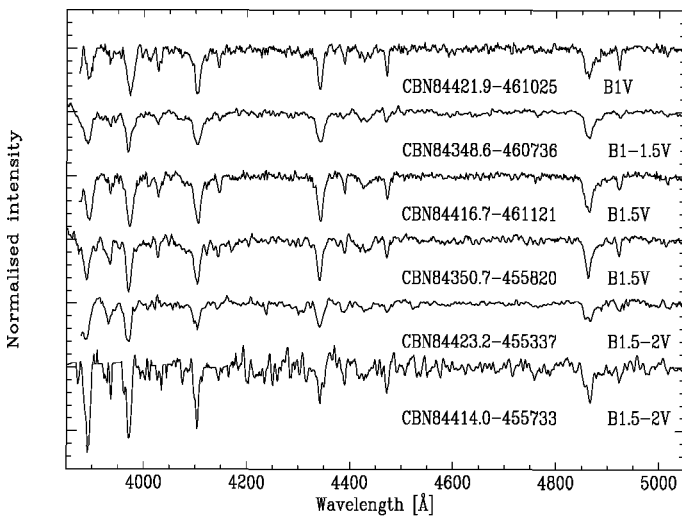


Fig. 8. This figure includes stars with B1-2V spectral types in the region of Bochum 7. 1.

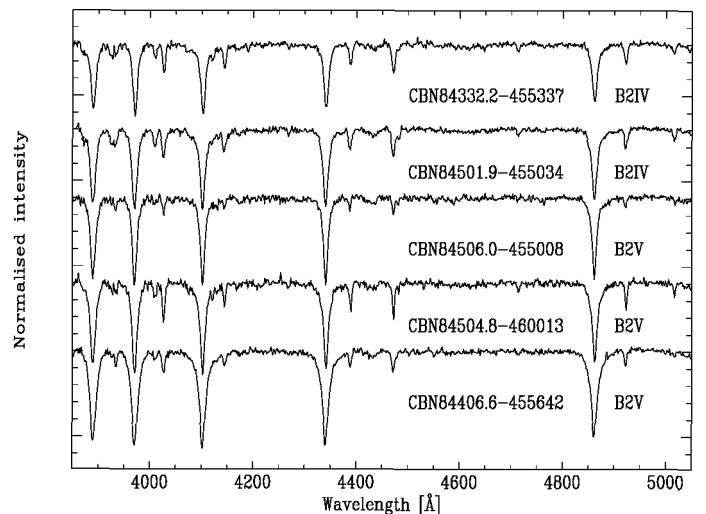


Fig. 11. This figure includes stars with B2IV-V spectral types in the region of Bochum 7. 1.

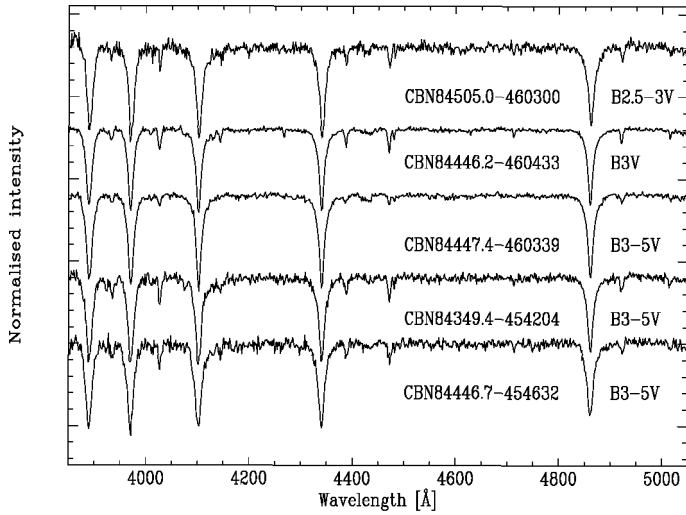


Fig. 12. This figure includes stars with B3-5V spectral types in the region of Bochum 7. 1.

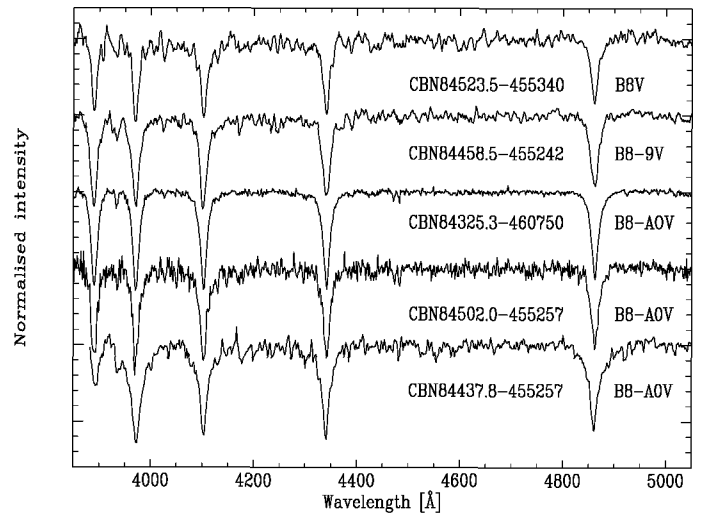


Fig. 15. This figure includes stars with B8-9V and A0V spectral types in the region of Bochum 7. 1.

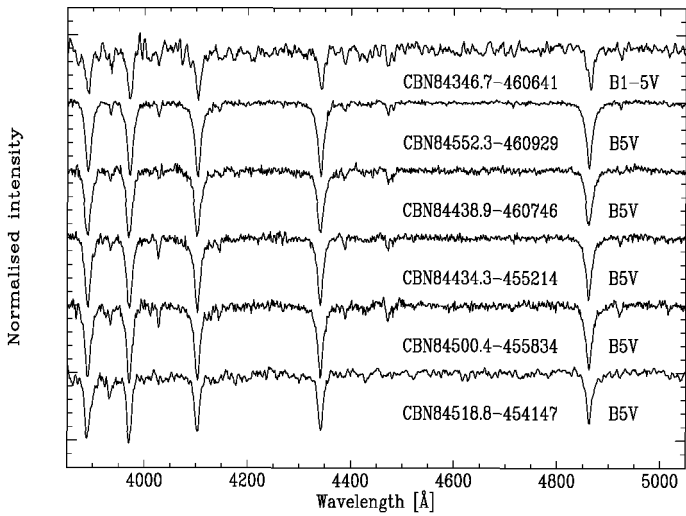


Fig. 13. This figure includes stars with B5V spectral types in the region of Bochum 7. 1.

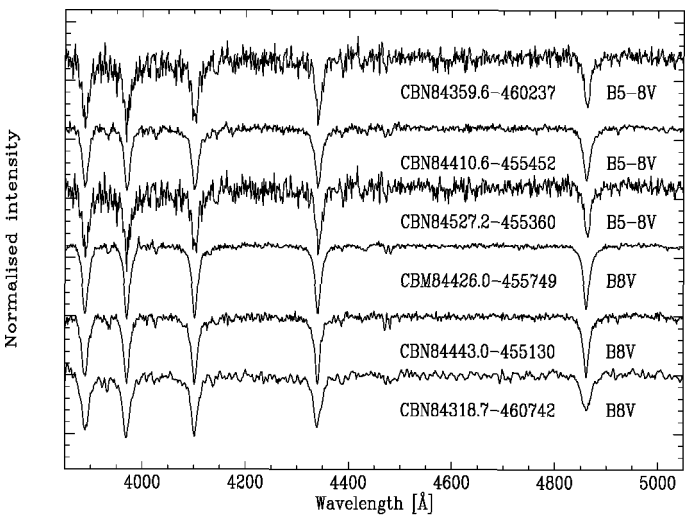


Fig. 14. This figure includes stars with B5-8V spectral types in the region of Bochum 7. 1.